

Sussing merger trees: the influence of the halo finder

Article (Published Version)

Avila, Santiago, Knebe, Alexander, Pearce, Frazer R, Schneider, Aurel, Srisawat, Chaichalit, Thomas, Peter, Behroozi, Peter, Elahi, Pascal J, Han, Jiaxin, Mao, Yao-Yuan, Onions, Julian, Rodriguez-Gomez, Vicente and Tweed, Dylan (2014) Sussing merger trees: the influence of the halo finder. *Monthly Notices of the Royal Astronomical Society*, 441 (4). pp. 3488-3501. ISSN 0035-8711

This version is available from Sussex Research Online: <http://sro.sussex.ac.uk/id/eprint/49217/>

This document is made available in accordance with publisher policies and may differ from the published version or from the version of record. If you wish to cite this item you are advised to consult the publisher's version. Please see the URL above for details on accessing the published version.

Copyright and reuse:

Sussex Research Online is a digital repository of the research output of the University.

Copyright and all moral rights to the version of the paper presented here belong to the individual author(s) and/or other copyright owners. To the extent reasonable and practicable, the material made available in SRO has been checked for eligibility before being made available.

Copies of full text items generally can be reproduced, displayed or performed and given to third parties in any format or medium for personal research or study, educational, or not-for-profit purposes without prior permission or charge, provided that the authors, title and full bibliographic details are credited, a hyperlink and/or URL is given for the original metadata page and the content is not changed in any way.

Impact of baryons on the cluster mass function and cosmological parameter determination

Sam J. Cusworth,¹★ Scott T. Kay,¹ Richard A. Battye¹ and Peter A. Thomas²

¹Jodrell Bank Centre for Astrophysics, School of Physics and Astronomy, The University of Manchester, Manchester M13 9PL, UK

²Astronomy Centre, Department of Physics and Astronomy, University of Sussex, Brighton BN1 9QH, UK

Accepted 2014 January 15. Received 2014 January 6; in original form 2013 September 15

ABSTRACT

Recent results by the Planck Collaboration have shown that cosmological parameters derived from the cosmic microwave background anisotropies and cluster number counts are in tension, with the latter preferring lower values of the matter density parameter, Ω_m , and power spectrum amplitude, σ_8 . Motivated by this, we investigate the extent to which the tension may be ameliorated once the effect of baryonic depletion on the cluster mass function is taken into account. We use the large-volume Millennium Gas simulations in our study, including one where the gas is pre-heated at high redshift and one where the gas is heated by stars and active galactic nuclei (in the latter, the self-gravity of the baryons and radiative cooling are omitted). In both cases, the cluster baryon fractions are in reasonably good agreement with the data at low redshift, showing significant depletion of baryons with respect to the cosmic mean. As a result, it is found that the cluster abundance in these simulations is around 15 per cent lower than the commonly adopted fit to dark matter simulations by Tinker et al. for the mass range $10^{14}–10^{14.5} h^{-1} M_\odot$. Ignoring this effect produces a significant artificial shift in cosmological parameters which can be expressed as $\Delta[\sigma_8(\Omega_m/0.27)^{0.38}] \simeq -0.03$ at $z = 0.17$ (the median redshift of the *Planck* cluster sample) for the feedback model. While this shift is not sufficient to fully explain the *Planck* discrepancy, it is clear that such an effect cannot be ignored in future precision measurements of cosmological parameters with clusters. Finally, we outline a simple, model-independent procedure that attempts to correct for the effect of baryonic depletion and show that it works if the baryon-dark matter back-reaction is negligible.

Key words: cosmological parameters – cosmology: theory.

1 INTRODUCTION

Clusters of galaxies map peaks in the cosmic density field and as such can be used to determine information about the Universe on the largest scales (Voit 2005; Allen, Evrard & Mantz 2011; Kravtsov & Borgani 2012). In particular, the abundance of clusters as a function of mass and redshift has been shown to be a particularly sensitive probe of cosmological parameters Ω_m and σ_8 , the matter density parameter and the linear rms matter fluctuation within a spherical top-hat of $8 h^{-1}$ Mpc radius, respectively¹ [Vikhlinin et al. 2009; Rozo et al. 2010; Reichardt et al. 2012; Hasselfield et al. 2013; Ade et al. (Planck Collaboration) 2013b]. In order to link the observed mass function of clusters to an underlying cosmology one must appeal to an analytic description of cluster abundance (Press & Schechter 1974; Bond et al. 1991; Sheth & Tormen 2002) or to one

of many numerical studies investigating dark matter halo formation, e.g. Jenkins et al. (2001), Tinker et al. (2008), Watson et al. (2013).² One of the most commonly adopted descriptions of cluster halo abundance is the Tinker et al. (2008) mass function (TMF hereafter).

The implicit assumption made in linking simulated dark matter halo masses with galaxy cluster masses is that the ratio of baryons to dark matter within clusters does not differ significantly from the cosmic value. This assumption, however, has been challenged by multiwavelength observations (Lin, Mohr & Stanford 2003; Giodini et al. 2009; Laganá et al. 2011). It has also been shown in N -body simulations that the pressure forces within baryonic gas are capable of segregating the distribution of collisional gas relative to pressureless dark matter (Navarro & White 1993), thereby changing the baryon fraction within clusters (Crain et al. 2007). Numerical studies have also shown that galaxy formation processes and non-gravitational heating can modify the baryon fraction (McCarthy

★ E-mail: cusworth@jb.man.ac.uk

¹ Throughout we express the Hubble parameter today as $H_0 = 100 h \text{ km s}^{-1} \text{ Mpc}^{-1}$.

² See Murray, Power & Robotham (2013) for a recent comparison of mass functions in the literature.

et al. 2011; Planck et al. 2013) and thereby the total mass within clusters (Stanek, Rudd & Evrard 2009).

The measurement of the cluster mass function from observations requires the calibration of an *observable-mass* ($X-M$) relation, where common observables, X , are X-ray luminosity, galaxy richness and Sunyaev–Zel’dovich (SZ) flux. Multiple ongoing observational surveys including the *Planck* mission [Ade et al. (Planck Collaboration) 2013b], the South Pole Telescope survey (Reichardt et al. 2012), Dark Energy Survey (The Dark Energy Survey Collaboration 2005) and the *XMM* Cluster Survey (XCS; Romer et al. 2001; Sahlén et al. 2009) have made the cosmological analysis of the galaxy cluster mass function one of their key scientific goals. In order to parametrize the systematic uncertainties in the measurement of a given scaling relation, e.g. of incorrectly assuming hydrostatic equilibrium in clusters, the mass bias parameter $b_{\text{hyd}} = 1 - M_{\text{true}}/M_X$ is commonly employed, where M_X is the mass inferred from observable X .

In the near future, large-volume observational surveys such as *eROSITA* (Pillepich, Porciani & Reiprich 2012), *Euclid* (Laureijs et al. 2011), the Large Synoptic Survey Telescope [Abell et al. (LSST Science Collaboration) 2009] and the proposed PRISM mission³ will detect a greater number of galaxy clusters than ever before. It is therefore of great importance that the cluster mass function is accurately calibrated against theoretical predictions (Reed et al. 2013).

Recent results from the *Planck* cluster survey [Ade et al. (Planck Collaboration) 2013b] have been found to be in tension with cosmological parameter determinations made using anisotropies in the cosmic microwave background [CMB; Ade et al. (Planck Collaboration) 2013c]. It has been argued that the discrepancy between the two measured values of σ_8 and Ω_m could be due, in part, to cluster biases and selection effects. Alternatively, it has been proposed that the influence of additional physical processes, such as the influence of massive neutrinos on the power spectrum, could lead to an underestimation in the mass function.

In this paper, we use large cosmological simulations with baryonic physics to investigate whether such tension can at least in part be explained by the effects of baryonic depletion in clusters. Such an effect, due to gas being expelled by feedback processes, produces a shift in the cluster mass function to lower abundance at fixed mass which if not accounted for in the cosmological analysis, leads to derived values for cosmological parameters (Ω_m and σ_8) that are systematically underestimated.

The remainder of the paper can be summarized as follows. In Section 2, we outline details of the simulations used and how the cluster samples were defined. In Section 3, our main results are presented, quantifying the effect of the baryon depletion on the mass function and its subsequent effect on the cosmological parameters Ω_m and σ_8 , before suggesting a simple corrective procedure. Finally, in Section 4, we discuss our results in the context of other work in the literature and draw conclusions.

2 SIMULATIONS AND HALO SELECTION

2.1 Millennium Gas simulations

We use results from three cosmological Millennium Gas simulations (MGS; Hartley et al. 2008; Stanek et al. 2009, 2010; Short et al. 2010; Young et al. 2011; Kay et al. 2012) and two dark matter-only

versions of the same volumes. The MGS are designed to include the dynamics of gas that were not present in the dark matter-only Millennium simulations. Each simulation in the suite is run with a different treatment of large-scale baryonic physics. We group these simulations into two ‘generations’, determined by the underlying cosmological model employed.

2.1.1 First generation: GO and PC models

In the *Gravitation Only* (GO) simulation, first described in Crain et al. (2007), baryonic gas is only permitted to change in entropy through shock heating. As a counterpoint to the adiabatic GO simulation, in the *Pre-heating & Cooling* (PC) simulation, described in Hartley et al. (2008), radiative cooling of gas was implemented (assuming a metallicity $Z = 0.3 Z_\odot$). Furthermore, in order to emulate the effects of high-redshift galaxy formation and reproduce the observed X-ray luminosity–temperature relation at $z \simeq 0$, the gas within the volume was uniformly heated to 200 keV cm² at $z = 4$.

The GO and PC simulations were run using the GADGET-2 code (Springel 2005) with the same cosmological model as the Millennium simulation (Springel et al. 2005). The parameters used were $\Omega_m = 0.25$, $\Omega_b = 0.045$, $h = 0.73$ and $\sigma_8 = 0.9$; consistent with the first year *Wilkinson Microwave Anisotropy Probe* results (WMAP1; Spergel et al. 2003). Because of computational constraints, the simulations were run with a slightly decreased mass resolution compared to the original Millennium run. A downgraded version of the Millennium initial conditions was used in the GO and PC simulations. At early times ($z > 3$), the gravitational softening length was fixed in comoving coordinates to $\epsilon = 100 h^{-1}$ kpc, whereas at late times ($z < 3$) the softening was fixed in physical coordinates to $\epsilon = 25 h^{-1}$ kpc. The particle masses were set to $m_{\text{dm}} = 1.4 \times 10^{10} h^{-1} M_\odot$ and $m_{\text{gas}} = 3.1 \times 10^9 h^{-1} M_\odot$ for the dark matter and gas, respectively. In both simulations, the dark matter was evolved self-consistently with the gas. As such the baryons influence the formation and growth rate of dark matter structures.

We compare the first generation MGS to a version of the original Millennium simulation (Springel et al. 2005) with the same initial conditions, mass resolution and gravitational softening lengths as the GO and PC models. We will refer to this simulation as DM1.

2.1.2 Second generation: FO model

In the *Feedback Only* (FO) simulation, the effects of stochastic active galactic nuclei (AGN) and supernovae feedback on the gas dynamics were inferred using the semi-analytic model of Guo et al. (2011). For full details regarding the treatment of the gas dynamics in FO; see Short, Thomas & Young (2013).⁴ The principal improvement of the FO simulation over the PC is that the baryonic feedback is better physically motivated, although radiative cooling is not included. One caveat of note is that in the FO model the baryonic contribution to the gravitational potential is ignored. In other words, the gas is evolved with zero gravitational mass and so there is no baryon–dark matter back-reaction.

The FO simulation was run using an updated version of the GADGET code, GADGET-3 at the resolution of the original Millennium simulation. Smaller softening lengths than the first generation simulations ($\epsilon = 37 h^{-1}$ kpc in comoving coordinates

³ <http://www.prism-mission.org/>

⁴ The MGS2-FO simulation described in Hilton et al. (2012) implemented the same physical model, albeit with a smaller simulated volume.

before $z \simeq 3$ and $\epsilon = 9.3 h^{-1} \text{ kpc}$ in physical coordinates there-after) were set. The masses of dark matter and gas particles were set to $7.8 \times 10^8 h^{-1} M_\odot$ and $3.1 \times 10^8 h^{-1} M_\odot$, respectively.

We will compare FO to an updated version of the dark matter-only Millennium simulation (DM2; Springel et al. 2005). Both simulations in the second generation used the same set of initial conditions, generated using second-order Lagrangian perturbation theory (Scoccimarro 1998). In addition, the FO and DM2 simulations were carried out using cosmological parameters consistent with the 7 yr WMAP results (WMAP7; Komatsu et al. 2011): $\Omega_m = 0.272$, $\Omega_b = 0.0455$, $h = 0.704$ and $\sigma_8 = 0.81$.

2.2 Cluster sample

Clusters were identified from the simulated density field using combinations of the Friends of Friends (FOF; Davis et al. 1985), SUBFIND (Springel et al. 2001) and spherical overdensity (Press & Schechter 1974; Lacey & Cole 1994) algorithms [see Knebe et al. (2011) for a review of halo-finding techniques]. In all of the analysis presented here, we consider clusters with mass $\mathcal{O}(10^{14} h^{-1} M_\odot)$. These clusters correspond to groups containing $>10^4$ particles.

In the first generation MGS (GO PC) and DM1 simulation, clusters were identified using the procedure outlined in Kay et al. (2012). Briefly, the dark matter particles were initially grouped using an FOF algorithm, with dimensionless linking length $b = 0.1$. The linking length parameter was chosen to be smaller than the canonical $b = 0.2$ in order to avoid the so-called overbridging problem whereby distinct, neighbouring haloes are linked together (Knebe et al. 2011). Next, the centre of each cluster was identified as the dark matter particle with the most negative gravitational potential energy. Finally, the bulk properties of the clusters (mass, radii, etc.) were calculated using the properties of all particles within spherical regions of overdensity $\Delta = \rho(<R_\Delta)/\rho_c(z)$ where $\rho_c(z) = (3H_0^2/8\pi G)E(z)^2$ is the cosmic critical density at redshift z and $E(z)^2 = \Omega_m(1+z)^3 + \Omega_\Lambda$. Throughout we take $\Delta = 500$, therefore cluster masses are defined

$$M_{500} = 500 \frac{4\pi}{3} R_{500}^3 \rho_c(z), \quad (1)$$

where R_{500} is the proper radius of the spherical overdensity. For reasons described in Section 3.1, we will consider clusters with $M_{500} = 10^{14} - 10^{14.5} h^{-1} M_\odot$. In the GO, PC and DM1 simulations there are 1016, 800 and 965 clusters, respectively, at $z = 0$.

In the second generation MGS (FO) and dark matter-only DM2 simulation, a similar procedure was implemented. First, an FOF algorithm was run with $b = 0.2$. In order to avoid the overbridging problem, SUBFIND was used to identify gravitationally bound structures within each FOF group. We then took the centre of the most massive substructure within the FOF groups to be the cluster centre. Finally, the bulk properties were calculated using the spherical overdensity algorithm. At $z = 0$, there are 707 and 830 clusters with M_{500} within the range of interest, for the FO and DM2 models, respectively. The different number of clusters found in DM1 and DM2 is mainly due to the fact that they utilize different cosmological models.

It should be noted that since the peak of the density field within an FOF group is considered to be the centre of a cluster in the analysis of both generations, the procedures used here are largely equivalent. The selection criteria in both generations exclude low-mass clusters whose centres lie within the R_{500} of more massive clusters, in line with other studies (Tinker et al. 2008).

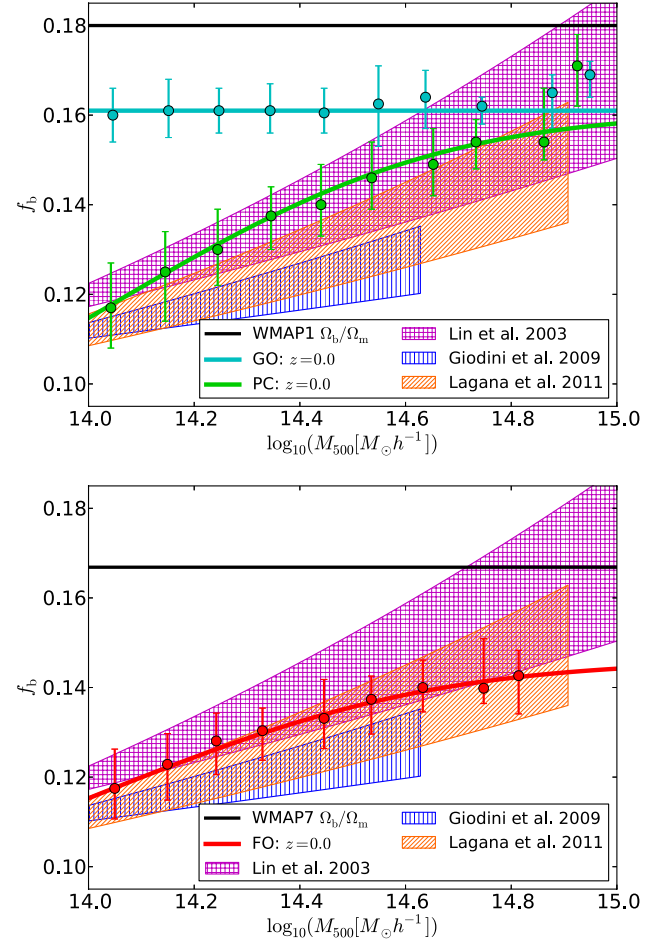


Figure 1. The baryon fraction as a function of M_{500} is shown for both generations of MGS. results from GO and PC are shown in the top panel (cyan and green, respectively) while results from FO are shown in red in the lower panel. The simulated data derived from simulation outputs at $z = 0$ are plotted, where bars indicate the 16th and 84th percentiles of the distribution and the coloured points show the median value within each mass bin. In the top panel, we plot the fits to the simulated data (GO and PC) from Young et al. (2011), and in the lower panel, we plot the fit computed for the FO simulation. We also plot the low-redshift observational bounds of Lin et al. (2003), Giodini et al. (2009) and Laganá et al. (2011) in the orange, purple and blue regions, respectively. In each panel, the cosmic mean, Ω_b/Ω_m , calculated using the corresponding WMAP 1/7 parameters is also shown in black.

2.3 Baryon fraction

In the simulated clusters, we define the baryon fraction

$$f_b = \frac{M_*(<R_{500}) + M_{\text{gas}}(<R_{500})}{M_{500}}, \quad (2)$$

where M_* and M_{gas} are the masses of stars and gas, respectively, within R_{500} .

The baryon fractions calculated from the clusters in the two generations of MGS are plotted in Fig. 1 as a function of cluster mass M_{500} . The baryon fraction for the GO and PC clusters, presented in Young et al. (2011), is also shown in the top panel of Fig. 1 (cyan and green curves, respectively). We follow Young et al. (2011) and fit the FO baryon fraction scaling relation to the function

$$\log_{10} f_b = \log_{10} f_0 + s \left[\mu - \frac{1}{4} \ln(1 + \exp(4\mu)) \right], \quad (3)$$

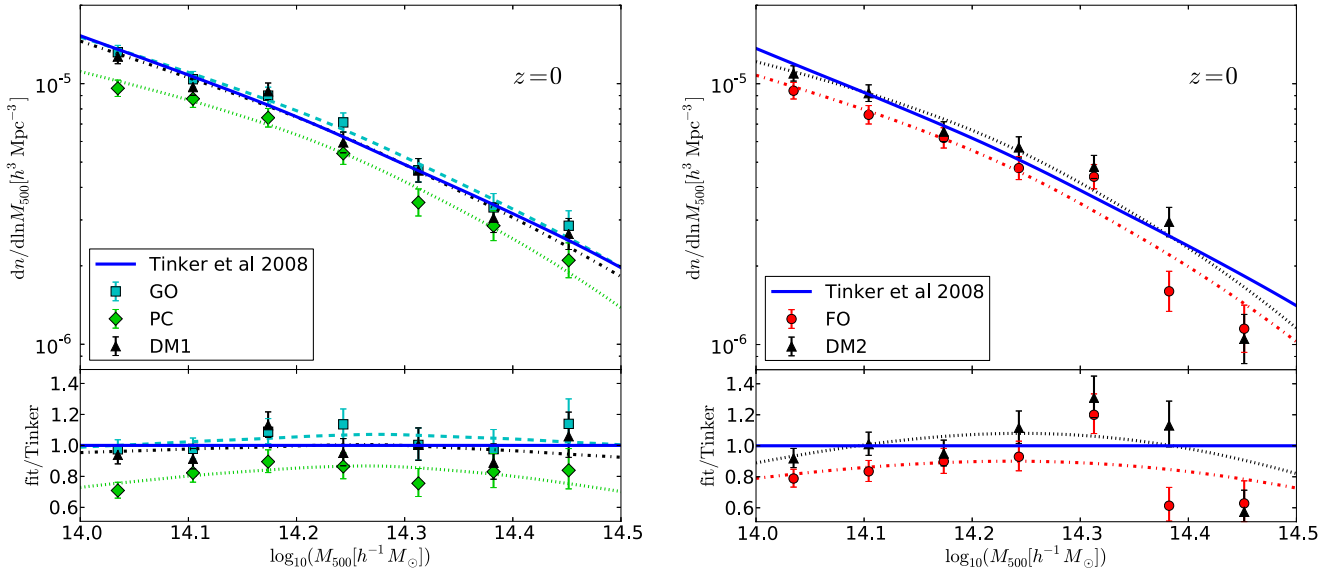


Figure 2. Top: differential mass functions plotted from the two generations of MGS. In the left-hand panel the GO and PC results are shown as cyan squares and green diamonds, respectively. In the right-hand panel the FO and DM2 mass functions are plotted as red circles and black triangles, respectively. Also plotted are fits to each cluster population computed by allowing the parameters in equation (6) to vary. In each panel, the TMF is also shown in solid blue. Lower: ratio of each mass function best fit with the TMF. Over the range plotted here the mean fit/Tinker values are $\simeq 0.82$, 0.86 for the PC and FO mass functions, respectively. For a similar comparison of GO and PC (using $\Delta = 200$) see fig. 4 of Stanek et al. (2009).

where $\mu = \log_{10}(M/M_{\text{piv}})$ is the mass variable scaled by a pivot mass M_{piv} and f_0 and s are two free parameters. The best-fitting parameters, for $\log_{10}(M_{\text{piv}} [h^{-1} M_{\odot}]) = 14.47$, were $f_0 = 0.146 \pm 0.001$ and $s = 0.204 \pm 0.008$. Errors quoted here were calculated using bootstrap resampling, keeping M_{piv} fixed.

In both panels of Fig. 1, we also plot the best fits to the observational data of Lin et al. (2003), Giodini et al. (2009) and Laganá et al. (2011) along with the associated uncertainties. Estimations of f_b within clusters require knowledge of bulk properties, such as mass and radius, measurements of the intracluster gas and observations of the stellar mass distribution. These observations are often made difficult by the contributions of intracluster light and fainter dwarf galaxies.

In non-radiative simulations, such as GO, the baryonic gas distribution within a cluster becomes more extended than the dark matter because it is able to gain energy in halo merger events (Crain et al. 2007). The resulting baryon fraction within R_{500} is therefore slightly reduced relative to the cosmic mean in a manner that is independent of halo mass. As a counterpoint, the baryon fraction in the PC and FO simulations is scale dependent. In the PC case, gas heated within a small halo is more likely to be ejected than gas in a halo with a deeper gravitational potential well. Similarly, the net effect of AGN and supernovae feedback is to eject more baryonic gas from lower mass clusters.

It is clear from Fig. 1 that the baryon fraction in $z \simeq 0$ galaxy clusters, both observed and simulated, is less than the cosmic mean Ω_m/Ω_b for the mass range plotted. There is excellent agreement between the simulated baryon fraction and the observational bounds plotted in both the PC and FO models.⁵ While the baryon fraction

within clusters can itself be used as a cosmological probe (Allen et al. 2011), here we use it as a test of the validity of the gas physics model employed in our simulations.

The baryonic depletion in the hydrodynamically simulated clusters leads directly to lower values of M_{500} relative to the dark matter-only counterparts. The effect of the depletion on the cluster mass function and subsequent cosmological parameter estimations is the subject of the following section.

3 RESULTS

3.1 Cluster mass function

Following Jenkins et al. (2001), we express the halo mass function as the logarithmic derivative of the number density, $n(M_{500})$, with respect to mass

$$\frac{dn}{d \ln M_{500}} = f(\sigma) \bar{\rho}_m(z) \frac{d \ln \sigma^{-1}}{d \ln M_{500}}, \quad (4)$$

where the variance of the density field within spheres of radius R [h^{-1} Mpc],

$$\sigma^2(R, z) = \frac{D^2(z)}{2\pi^2} \int P(k) W^2(kR) k^2 dk, \quad (5)$$

$P(k)$ is the linear matter power spectrum, $D(z)$ is the linear growth factor and $W(x) = 3(\sin x - x \cos x)/x^3$ is the Fourier transform of the real-space top-hat filter. The function $f(\sigma)$ is independent of cosmological parameters by design. Recent studies (Tinker et al. 2008; Watson et al. 2013) have taken the parametrization

$$f(\sigma) = A \left[\left(\frac{\beta}{\sigma} \right)^\alpha + 1 \right] \exp \left(-\frac{c}{\sigma^2} \right), \quad (6)$$

and computing the constants A , α , β , c from their respective dark matter-only cosmological simulations for a range of z and Δ .

Fig. 2 shows mass functions computed from the cluster distributions in the MGS at $z = 0$, where the mass bins were spaced with

⁵ We note that the baryon fraction in low mass simulated clusters ($M_{500} \lesssim 10^{14} h^{-1} M_{\odot}$; not shown) was found to be significantly lower than the cosmic mean ($f_b \simeq 0.1$) in agreement with Lin et al. (2003) and simulations which include AGN feedback (e.g. McCarthy et al. 2011; Planelles et al. 2013).

$\Delta \ln M_{500} = 0.16$ and the position of each bin was taken to be the mid-point.⁶ We also plot the appropriate TMF in both panels, where $P(k)$ was calculated using the publicly available code CAMB⁷ (Lewis, Challinor & Lasenby 2000). It is evident from the right-hand panel that the TMF is broadly consistent with the dark matter-only simulation for $14 \lesssim \log_{10}(M_{500}[h^{-1} M_{\odot}]) \lesssim 14.5$. Above this mass, Poisson noise due to rare objects starts to become significant. The agreement between the DM1 and DM2 mass functions and the TMF is within the 5 per cent statistical errors of the TMF fitting at the low-mass end of the mass function. For confirmation of the agreement between DM1/DM2 and the TMF, see Fig. 3.

3.2 Impact of baryons on mass function

As discussed in Stanek et al. (2009), the clusters in the PC simulation showed a systematic suppression relative to both the GO clusters and the TMF. We demonstrate this effect again in the left-hand panel of Fig. 2 for $\Delta = 500$. One can also see from the right-hand panel of Fig. 2 that the mass function computed from FO is also offset from both the data of DM2 and the TMF.

We note that there is agreement between the GO mass function and the TMF. In this model, the mass-independent baryonic depletion, detailed in Section 2.3, is sufficiently mild that the expected underlying dark matter-only (DM1) mass function is recovered. In both the PC and FO mass functions, the larger relative offset is a consequence of the lower, mass-dependent, cluster baryon fraction resulting from gas ejection processes.

One can parametrize the deviation from the TMF by generalizing the mass bias parameter described in Section 1 to include the effects of baryon depletion. Since we know the true masses of our simulated clusters, we will ignore the complexities of hydrostatic bias in cluster observations. Instead, we define the baryonic depletion bias $b_{\text{dep}} = 1 - M_{\text{DM}}/M_{\text{hyd}}$, where M_{DM} and M_{hyd} are the masses of a cluster in dark matter-only and hydrodynamic simulations, respectively. We found that rescaling the mass variable, M_{500} , in the DM2 mass function at $z = 0$ ($z = 0.17$) by $1 - b_{\text{dep}} = 0.9$ (0.93) brought the FO and DM2 curves into closer agreement. We note, however, a noticeable difference in shape between the FO and adjusted DM2 mass functions was evident. While the zeroth-order effect of the baryons on the mass function is to shift it relative to the dark matter-only mass function by 10 per cent (7 per cent), we argue that it is insufficient for precision cosmology. Further, arguments of this type do not account for changes in R_{500} in a consistent manner.

We will return to the problem of baryonic influence on the mass function in Section 3.4 and discuss the effect of back-reaction of baryons on the dark matter. It will be shown that the primary reason for the change in mass is ejection of baryons from clusters.

3.3 Effect on cosmological parameters

As outlined in the introductory section, the primary function of the TMF is to link measurements of the halo mass function to the Ω_m and σ_8 parameters of the underlying cosmology. By using the TMF to constrain cosmology from galaxy cluster measurements, it is assumed that is that the gas content of clusters traces the dark matter component. Given that we have demonstrated that simulating

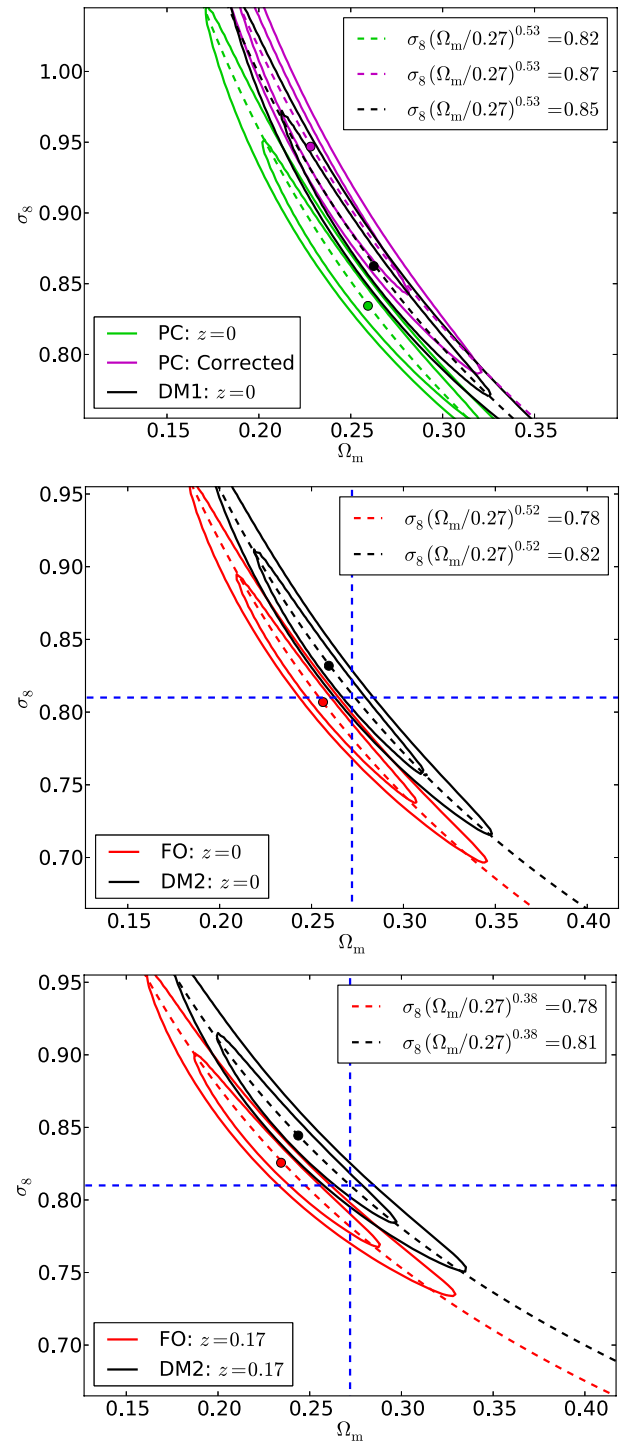


Figure 3. Likelihood contours computed from the simulated cluster mass functions at $z = 0$ (top and middle) and $z = 0.17$ (lower) assuming the TMF (see Fig. 2). In the top panel, cyan contours were computed using the GO mass function and the green contours were calculated using the results of the PC simulation. Similarly in the middle and lower panels, the DM2 and FO likelihood contours are shown in black and red, respectively. Also shown are the lines of best fit describing the degeneracy between Ω_m and σ_8 for each generation. In order to directly evaluate the shift in the degeneracy, we enforce the PC and GO power-law indices to be that of the DM1 and the FO power-law index to be that of the DM2. The points of maximum likelihood are shown as coloured dots. The discrepancy between the PC and FO distributions and the fiducial values of Ω_m and σ_8 (indicated by the blue dashed lines) is the key result of our investigations.

⁶ We have confirmed that the conclusions of this paper are not sensitive to either taking the bin mid-point, rather than the mean or median, nor the logarithmic width of the bins.

⁷ <http://camb.info>

the baryonic content within clusters suppresses cluster abundance at fixed mass relative to the dark matter-only result (particularly in the PC and FO cases), we now investigate the impact of this result on estimations of cosmological parameters.

We use the simulated cluster mass functions described in the previous sections (see Fig. 2) as our mock data. Taking the TMF as our assumed model, we computed likelihood distributions for each of the simulated populations on a regular grid.

At each point in Ω_m, σ_8 space, the Cash statistic (Cash 1979)

$$C = -2 \sum_{k=1}^{N_{\text{tot}}} \ln \mathcal{P}(N_k | n_k) \\ = -2 \sum_{k=1}^{N_{\text{tot}}} (N_k \ln(n_k) - n_k - \ln(N_k!)), \quad (7)$$

was calculated, where $\mathcal{P}(N_k | n_k)$ is the probability of finding N_k clusters in a bin given a number n_k predicted by the model. We then used the fact that ΔC is distributed as $\Delta\chi^2$ with two degrees of freedom (Press et al. 1992).

As in Section 3.1, we used CAMB to calculate $P(k)$ and hence the model $dn/d\ln M_{500}$ through equation (4). Throughout we assumed the values of other cosmological parameters were known since they do not contribute significantly to the variance in the mass function measurement (Murray et al. 2013).

The likelihood contours calculated using the MGS clusters are shown in Fig. 3. As before, in this analysis we conservatively used clusters with $14 < \log_{10}(M_{500}[h^{-1} M_\odot]) < 14.5$. We show the earlier epoch since it is the median z of the 2013 *Planck* SZ high S/N catalogue [Ade et al. (Planck Collaboration) 2013a]. The lines of degeneracy between σ_8 and Ω_m are also shown in Fig. 3.

In the upper panel of Fig. 3, the offset between the GO and PC likelihood contours is clear. While the peak of the $z = 0$ PC distribution is offset along the degeneracy, the movement of the degeneracy itself is the crucial characteristic.

The likelihood distribution contours resulting from the FO and DM2 mass functions are shown in right-hand panel of Fig. 3. The change in the FO/DM2 power-law indices of the degeneracy between the middle and bottom panels is due to the redshift dependence of the mass function. As the mass function evolves with redshift, it enables one, in principle, to break the Ω_m - σ_8 degeneracy with multiredshift observations. The FO contours are clearly shifted relative to the dark matter-only simulation at both epochs. We quantify this shift as $\Delta[\sigma_8(\Omega_m/0.27)^{\gamma_z}] = \Delta B_z$, where γ_z and ΔB_z are constants, by fitting a power-law relation to the likelihood contours. Over the redshifts of interest ($z = 0 \rightarrow 0.17$), γ_z varies from 0.52 to 0.38 whereas ΔB_z remains $\simeq -0.03$. The discrepancy between the *Planck* CMB and cluster count measurements can be described as $\Delta[\sigma_8(\Omega_m/0.27)^{0.3}] \simeq -0.08$. It should be noted that the widths of the error contours (but crucially not the offset ΔB_z) in Fig. 3 reflect the size of the simulation volume (and hence the number of clusters) rather than any particular observational survey. Though we urge caution when comparing simulated snapshots and observations, our calculations demonstrate that the effects of baryonic depletion in clusters is non-negligible in this context.

3.4 Correcting for baryonic physics in galaxy clusters

As we have demonstrated in the previous sections, the mass of a given cluster in a hydrodynamic simulation is not equal to the mass of the same object in a dark matter-only simulation. We now outline and test a method for ‘correcting’ a baryon influenced cluster mass

function in order to enable one to use the TMF (or similar) for cosmological parameter determinations.

Our proposed three-step methodology is as follows:

(i) calculate the *dark matter* mass profile of each cluster, $M(< r)$, using knowledge of the total density profile and removing the stellar and gas components;

(ii) supplement the dark matter mass with baryons such that the baryon fraction is equal to the cosmic value everywhere in the cluster, i.e. estimate the mass profile $M_{\text{est}}(< r) = M_{\text{DM}}(< r)/(1 - \Omega_b/\Omega_m)$;

(iii) Recalculate R_{500} and $M_{\text{est}, 500}$ using the new mass profile.

Observationally, it is too expensive to calculate the dark matter mass profile of every cluster in a survey. This would require, for example, high-quality X-ray data allowing the estimation of total density and temperature profiles, or weak lensing data with sufficiently high density of background sources. Additionally, the mass distribution of gas (using X-ray data) and stars (including any additional diffuse component) would also be required. In practice, a mass-observable relation (ideally with minimal scatter) is calibrated for a smaller number of clusters and the observable used as a mass *proxy* for the full sample (e.g. Arnaud, Pointecouteau & Pratt 2007). The practice of mass proxy calibration is common to all cluster surveys including the *Planck* analysis [Ade et al. (Planck Collaboration) 2013b]. In this case, the simple procedure outlined above could be applied to re-calibrate the mass-observable relation for cosmological purposes (we leave feasibility of such a procedure to future study).

The advantage of this procedure (rather than attempting to correct the theoretical mass function) is that it is empirical, relying on only the observational data and not assuming any theoretical model for how the baryons affect the total cluster mass. However, it does assume that the baryonic processes do not significantly influence the underlying dark matter mass profile of clusters.

We test our methodology using the clusters in the GO, PC and FO simulations. In Fig. 4, we plot the ratio of the mass of individual clusters in the hydrodynamical simulations and the dark matter-only simulations. We also plot the same mass ratio after the correction procedure was applied to the hydrodynamic cluster masses. The halo structures in PC were matched with their DM1 counterparts by considering the dark matter-only haloes within $0.5R_{500}$ of the cluster centre. A match was found for around 98 per cent of the PC clusters. In the FO case, each cluster was mapped directly to the equivalent DM2 halo using the data from the SUBFIND analysis of the dark matter distribution. In the absence of baryon-dark matter back-reaction, we expect the corrected distributions to be centred at unity.

In the FO simulation, the dark matter particles are not gravitationally influenced by the baryons and so no back-reaction is possible (Short et al. 2013). This fact is demonstrated by the purple distribution in the right-hand panel of Fig. 4 being centred around unity. The small amount of scatter is due to the error in calculating M_{500} from cluster profiles rather than the actual particles themselves. As a further test of our method, we confirm in the right-hand panel of Fig. 5 that the corrected FO mass function reproduces the DM2 mass function.

The GO and PC implementations, on the other hand, do evolve the dark matter and gas consistently. As shown in the upper and middle panels of Fig. 4, the median mass ratios for the GO and PC clusters are 1.046 ($\sigma_{\text{GO}} = {}^{+0.039}_{-0.052}$) and 0.949 ($\sigma_{\text{PC}} = {}^{+0.048}_{-0.045}$), respectively; whereas after correction procedure the median values are 1.020 and 1.053 with scatter $\sigma_{\text{GO}} = {}^{+0.033}_{-0.048}$ and $\sigma_{\text{PC}} = {}^{+0.047}_{-0.054}$. The

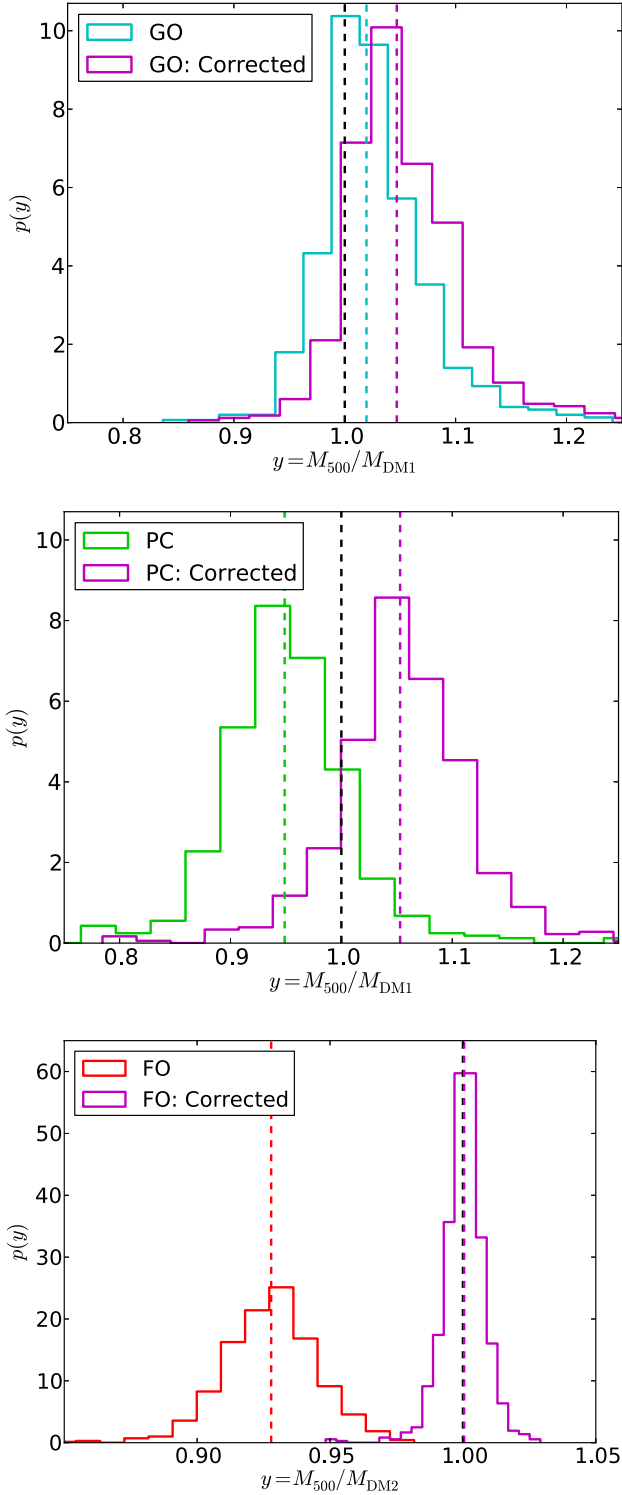


Figure 4. Distributions of total mass ratios, $y = M_{500}/M_{DM1/2}$, where M_{500} is the total mass of a cluster in the hydrodynamical simulation (GO, PC and FO are shown in cyan, green and red, respectively) and $M_{DM1/2}$ is the mass of the same cluster in the appropriate dark matter-only simulation. We also plot the ratio of the cluster masses after correcting the hydrodynamical mass as outlined in Section 3.4. In the FO model, the correction is near exact by construction; the ± 0.006 scatter about $y = 1.001$ is due to numerical error in recomputing R_{500} from the $M_{\text{est}}(<r)$ profile rather than the particle distribution. In the GO and PC models, where baryons can influence the dark matter density profile, the corrected halo masses do not exactly match the cluster masses in DM1.

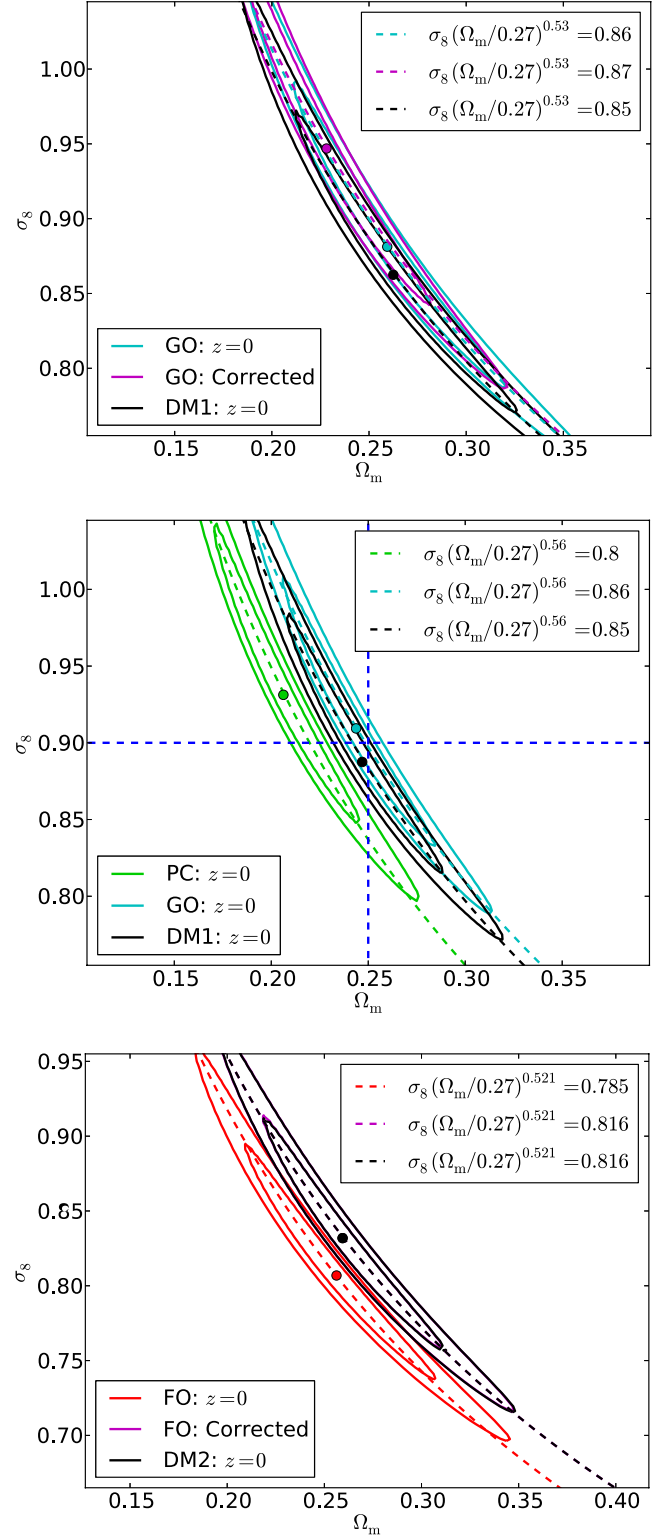


Figure 5. Likelihood contours calculated from $z = 0$ mass functions in the same manner as those in Fig. 3. ‘Corrected’ distributions were derived from the cluster mass function after applying the procedure to each cluster. As before, the results from the GO, PC and FO simulations are shown in cyan, green and red, respectively. The likelihood distributions calculated using the corrected mass functions are shown in purple. Additionally, the distributions obtained using the DM1/DM2 simulations are shown in black. The power-law index of the fitting is fixed to the value of the dark matter-only degeneracy.

quoted values of scatter were calculated from the 16th and 84th percentiles of the distributions. The fact that the corrected distributions are not centred on unity with minimal scatter is a reflection of the baryon-dark matter back-reaction.⁸ This irreducible effect is the central limitation of our method. Without a greater understanding of the effects of baryons on the gravitational potential, the procedure may be unable to recover the abundance of clusters to within around 5 per cent.

In Fig. 5, we compute likelihood contours from the GO, PC and FO mass functions before and after applying the above correction procedure. Note that because of the mass limit made on the PC cluster catalogue and the fact that the correction invariably increased the mass of clusters, we only consider clusters with $\log_{10}(M_{500}/h^{-1}M_{\odot}) > 14.1$. This additional condition decreases the number of clusters in the mass function and therefore increases the width of the contours shown in the left-hand panel of Fig. 5 as well as moving the contours along the σ_8 - Ω_m degeneracy. In the lower panel of Fig. 5, the corrected FO contours map almost directly on to the DM2 contours. This excellent agreement is due to the fact that, in the FO model, baryons do not influence the dark matter mass profile: scaling the dark matter mass profile by $(1 - \Omega_b/\Omega_m)^{-1}$ recovers the DM2 mass distribution by design.

As expected from the offset in the GO/DM1 and PC/DM1 cluster mass ratio distributions shown in Fig. 4, the contours from the corrected mass functions do not match the DM1 contours. In fact, the correction procedure does not appear to improve agreement between the GO/PC contours and the DM1. The fact that the input cosmology is not recovered by this correction procedure is a demonstration that the baryons do have a significant influence on the shape of the dark matter mass profile in these models.

4 DISCUSSION

In this paper, we have discussed the impact that baryonic physics can have on the observed cluster mass function. Although further study is required in order to fully model the gas physics in clusters, we have shown that the baryon fraction measured in MGS is broadly consistent with observations. We argue that since the baryon fraction is similar to that observed in other simulated clusters (Planelles et al. 2013; Sembolini et al. 2013), the suppression in the mass function shown in Fig. 2 is generic to realistic baryonic treatments.

In contrast to our findings, an ~ 7 per cent overabundance of clusters relative to dark matter only haloes was reported by Cui et al. (2012). In their simulations, efficient radiative cooling of gas ensures that the hydrodynamically simulated clusters are more concentrated than their dark matter-only counterparts and therefore have larger values of M_{500} . We reason that the lack of AGN feedback or early heating in their simulations allowed clusters to retain their baryon content and thereby allowed the mass of a given cluster to increase relative to its dark matter-only counterpart. Due to the functional form of the halo mass function, a shift in mass of this nature would result in an increase in the number of clusters of a fixed mass. As argued in van Daalen et al. (2011), a realistic treatment of gas dynamics results in a decrease in the matter power spectrum relative to dark matter-only simulations on scales $1 \lesssim k \lesssim 10 h \text{ Mpc}^{-1}$. It therefore follows that a relative underabundance of clusters would be expected.

We have shown that assuming the TMF leads to an incorrect measurement of the σ_8 - Ω_m degeneracy by $\Delta[\sigma_8(\Omega_m/0.27)^{0.38}] \approx$

-0.03 at $z = 0.17$ when considering realistic clusters with $14 < \log_{10}(M_{500}[h^{-1}M_{\odot}]) < 14.5$. The discrepancy we describe here is not specific to the TMF. We confirmed that using the Watson et al. (2013) fit as the assumed model, instead of the TMF, produced a similar offset between the derived dark matter-only and hydrodynamic likelihood distributions.

The analysis of Balaguera-Antolínez & Porciani (2013) came to similar conclusions as presented in this work, though through different means. In that study, predictions regarding the observed mass function were made using the TMF and assuming the form of $f_b(M_{500})$ from Laganá et al. (2011). By creating mock catalogues, the above authors showed a systematic shift of the similar order and sense as that shown in Fig. 3. Recently however, Martizzi et al. (2013) extended the Balaguera-Antolínez & Porciani methodology using the f_b - M_{500} relation derived from their set of high-resolution cluster resimulations. They concluded that the mass function should be boosted by the effects of baryonic physics because, in contrast with the observational data shown in Fig. 1, the baryon fraction they use is higher than the cosmic mean over the mass range. Our investigations differ from the above methodology in that we make no assumption regarding the functional form of the cosmic mass function or the f_b - M_{500} relation at the run-time of our simulations. Further, we have shown the Ω_m - σ_8 degeneracy offset is present in two very physically distinct scenarios (PC and FO).

In future *Planck* analyses, where lower mass clusters are studied, the influence of baryonic physics on cluster masses will have to be considered and accounted for. Further, our calculations are applicable to any cosmological survey with clusters of mass $\lesssim 10^{14.5} h^{-1} M_{\odot}$ (particularly XCS) and provide qualitative information on what might happen at larger masses, although we leave this for future investigations.

In our final section, we proposed and tested a model-independent procedure designed to recover the results of dark matter-only simulations from measurements of clusters with baryonic effects. The correction procedure was demonstrated to work well in the case where baryon-dark matter gravitational interaction was neglected (FO model). However in simulations in which baryons significantly contribute to the gravitational potential (GO/PC), the procedure was deemed to be insufficient. We conclude that further modelling of baryonic physics in clusters is required in order to ensure that future cluster surveys are able to make unbiased constraints on cosmological parameters.

ACKNOWLEDGEMENTS

We thank the referee, Joop Schaye, for his very helpful comments which have improved the quality of the manuscript. We also thank Frazer Pearce for providing the data from the DM1 simulation. This work used the DiRAC Data Centric system at Durham University, operated by the Institute for Computational Cosmology on behalf of the STFC DiRAC HPC Facility (www.dirac.ac.uk). This equipment was funded by BIS National E-infrastructure capital grant ST/K00042X/1, STFC capital grant ST/H008519/1, and STFC DiRAC Operations grant ST/K003267/1 and Durham University. DiRAC is part of the National E-Infrastructure. In addition, some of the simulations used in this paper were performed at the University of Nottingham HPC Facility. SJC is supported by an STFC quota studentship and the University of Manchester President's Doctorial Scholar Award. PAT acknowledges support from the Science and Technology Facilities Council (grant number ST/I000976/1).

⁸ We also checked that at lower overdensity ($\Delta = 200$), the difference between DM1 and GO masses was smaller than at $\Delta = 500$.

REFERENCES

- Abell P. A. et al. (LSST Science Collaboration), 2009, preprint ([arXiv:0912.0201](https://arxiv.org/abs/0912.0201))
- Ade P. A. R. et al. (Planck Collaboration), 2013a, preprint ([arXiv:1303.5089](https://arxiv.org/abs/1303.5089))
- Ade P. A. R. et al. (Planck Collaboration), 2013b, preprint ([arXiv:1303.5080](https://arxiv.org/abs/1303.5080))
- Ade P. A. R. et al. (Planck Collaboration), 2013c, preprint ([arXiv:1303.5076](https://arxiv.org/abs/1303.5076))
- Allen S. W., Evrard A. E., Mantz A. B., 2011, *ARA&A*, 49, 409
- Arnaud M., Pointecouteau E., Pratt G. W., 2007, *A&A*, 474, L37
- Balaguera-Antolínez A., Porciani C., 2013, *J. Cosmol. Astropart. Phys.*, 2013, 022
- Bond J. R., Cole S., Efstathiou G., Kaiser N., 1991, *ApJ*, 379, 440
- Cash W., 1979, *ApJ*, 228, 939
- Crain R. A., Eke V. R., Frenk C. S., Jenkins A., McCarthy I. G., Navarro J. F., Pearce F. R., 2007, *MNRAS*, 377, 41
- Cui W., Borgani S., Dolag K., Murante G., Tornatore L., 2012, *MNRAS*, 423, 2279
- Davis M., Efstathiou G., Frenk C. S., White S. D. M., 1985, *ApJ*, 292, 371
- Giodini S. et al., 2009, *ApJ*, 703, 982
- Guo Q. et al., 2011, *MNRAS*, 413, 101
- Hartley W., Gazzola L., Pearce F., Kay S., Thomas P., 2008, *MNRAS*, 386, 2015
- Hasselfield M. et al., 2013, *J. Cosmol. Astropart. Phys.*, 07, 008
- Hilton M. et al., 2012, *MNRAS*, 424, 2086
- Jenkins A., Frenk C. S., White S. D. M., Colberg J. M., Cole S., Evrard A. E., Couchman H. M. P., Yoshida N., 2001, *MNRAS*, 321, 372
- Kay S. T., Peel M. W., Short C. J., Thomas P. A., Young O. E., Battye R. A., Liddle A. R., Pearce F. R., 2012, *MNRAS*, 422, 1999
- Knebe A. et al., 2011, *MNRAS*, 415, 2293
- Komatsu E. et al., 2011, *ApJS*, 192, 18
- Kravtsov A. V., Borgani S., 2012, *ARA&A*, 50, 353
- Lacey C., Cole S., 1994, *MNRAS*, 271, 676
- Laganá T. F., Zhang Y.-Y., Reiprich T. H., Schneider P., 2011, *ApJ*, 743, 13
- Laureijs R. et al., 2011, preprint ([arXiv:1110.3193](https://arxiv.org/abs/1110.3193))
- Lewis A., Challinor A., Lasenby A., 2000, *ApJ*, 538, 473
- Lin Y., Mohr J. J., Stanford S. A., 2003, *ApJ*, 591, 749
- McCarthy I. G., Schaye J., Bower R. G., Ponman T. J., Booth C. M., Dalla Vecchia C., Springel V., 2011, *MNRAS*, 412, 1965
- Martizzi D., Mohammed I., Teyssier R., Moore B., 2013, preprint ([arXiv:1307.6002](https://arxiv.org/abs/1307.6002))
- Murray S., Power C., Robotham A., 2013, *MNRAS*, 434, L61
- Navarro J. F., White S. D. M., 1993, *MNRAS*, 265, 271
- Pillepich A., Porciani C., Reiprich T. H., 2012, *MNRAS*, 422, 44
- Planelles S., Borgani S., Dolag K., Ettori S., Fabjan D., Murante G., Tornatore L., 2013, *MNRAS*, 431, 1487
- Press W. H., Schechter P., 1974, *ApJ*, 187, 425
- Press W., Teukolsky S., Vetterling W., Flannery B., 1992, *Numerical Recipes in C*. Cambridge Univ. Press, Cambridge
- Reed D. S., Smith R. E., Potter D., Schneider A., Stadel J., Moore B., 2013, *MNRAS*, 431, 1866
- Reichardt C. L. et al., 2012, *ApJ*, 755, 70
- Romer K., Viana P., Liddle A., Mann R., 2001, *ApJ*, 547, 594
- Rozo E. et al., 2010, *ApJ*, 708, 645
- Sahlén M. et al., 2009, *MNRAS*, 397, 577
- Scoccimarro R., 1998, *MNRAS*, 299, 1097
- Sembolini F., Yepes G., De Petris M., Gottlöber S., Lamagna L., Comis B., 2013, *MNRAS*, 429, 323
- Sheth R. K., Tormen G., 2002, *MNRAS*, 329, 61
- Short C. J. J., Thomas P. A. A., Young O. E. E., Pearce F. R. R., Jenkins A., Muanwong O., 2010, *MNRAS*, 408, 2213
- Short C., Thomas P., Young O., 2013, *MNRAS*, 428, 1225
- Spergel D. N. et al., 2003, *ApJS*, 148, 175
- Springel V., 2005, *MNRAS*, 364, 1105
- Springel V., White S. D. M., Tormen G., Kauffmann G., 2001, *MNRAS*, 328, 726
- Springel V. et al., 2005, *Nature*, 435, 629
- Stanek R., Rudd D., Evrard A., 2009, *MNRAS*, 394, L11
- Stanek R., Rasia E., Evrard A. E., Pearce F., Gazzola L., 2010, *ApJ*, 715, 1508
- The Dark Energy Survey Collaboration, 2005, preprint ([astro-ph/0510346](https://arxiv.org/abs/astro-ph/0510346))
- Tinker J., Kravtsov A. V., Klypin A., Abazajian K., Warren M., Yepes G., Gottlöber S., Holz D. E., 2008, *ApJ*, 688, 709
- van Daalen M. P., Schaye J., Booth C. M., Dalla Vecchia C., 2011, *MNRAS*, 415, 3649
- Vikhlinin A. et al., 2009, *ApJ*, 692, 1060
- Voit G., 2005, *Rev. Mod. Phys.*, 77, 207
- Watson W. A., Iliev I. T., D'Aloisio A., Knebe A., Shapiro P. R., Yepes G., 2013, *MNRAS*, 433, 1230
- Young O. E., Thomas P. A., Short C. J., Pearce F., 2011, *MNRAS*, 413, 691

This paper has been typeset from a \LaTeX file prepared by the author.